Gamma-rays, cosmic-rays, infrared light and the intergalactic magnetic fields

Alexander Kusenko
(UCLA and Kavli IPMU)

NedFest, UCLA, August 25, 2017

Celebrating

the many areas of science in which

Ned

has made important lasting contributions

How to tell the moods of Ned Wright during your talk

intereste
d

bored

skeptical

enthusiastic

puzzled

excited

Quiz question: this image show Ned...

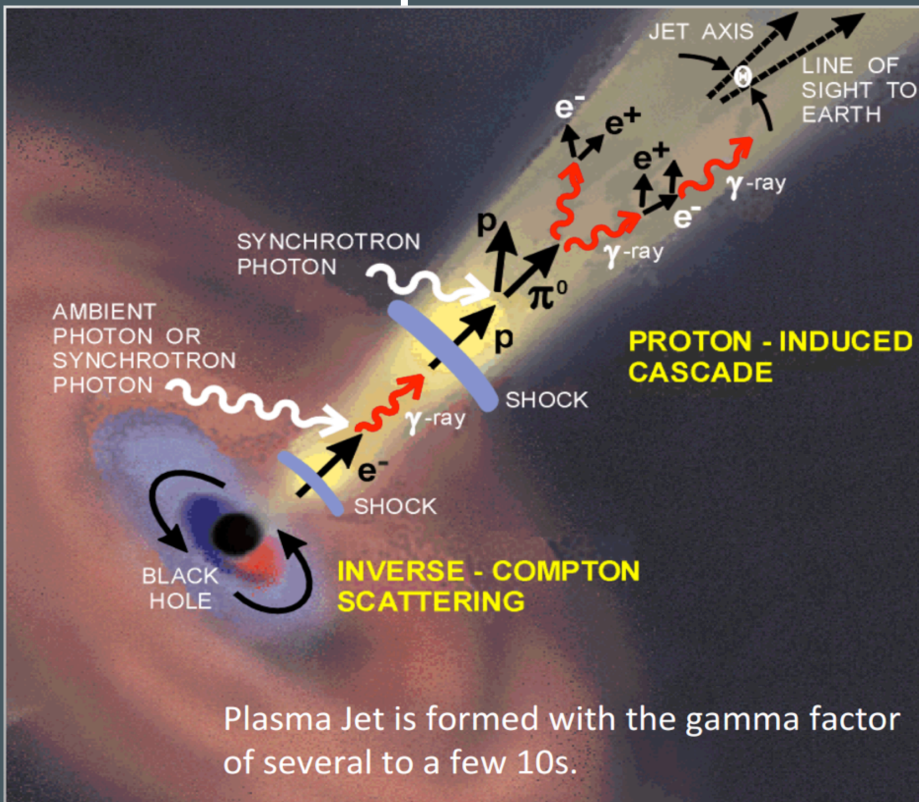
- ❑ Giving a presentation
- ❑ Teaching
- ❑ Flying
- ❑ Sailing
- ❑ Battling zombies
- ❑ None of the above

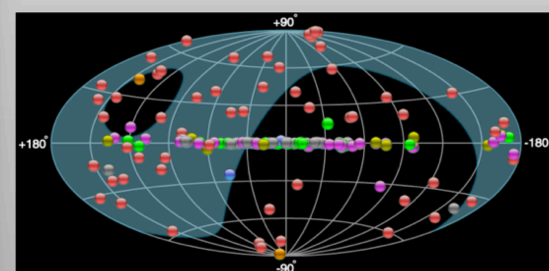
Correct answer: sailing

Quiz: find Ned in this image

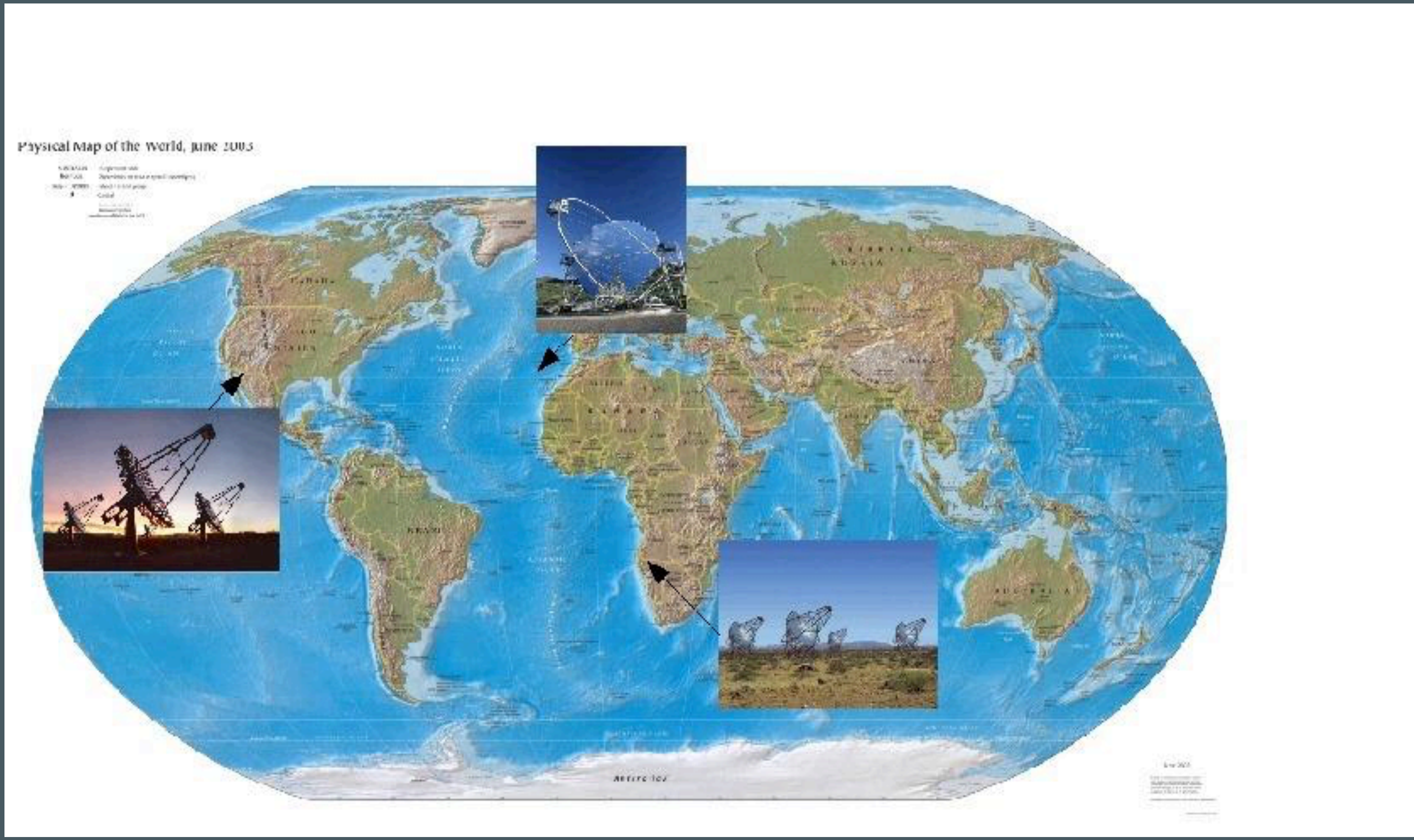
Quiz: find Ned in this image

Blazars: supermassive black holes with a jet

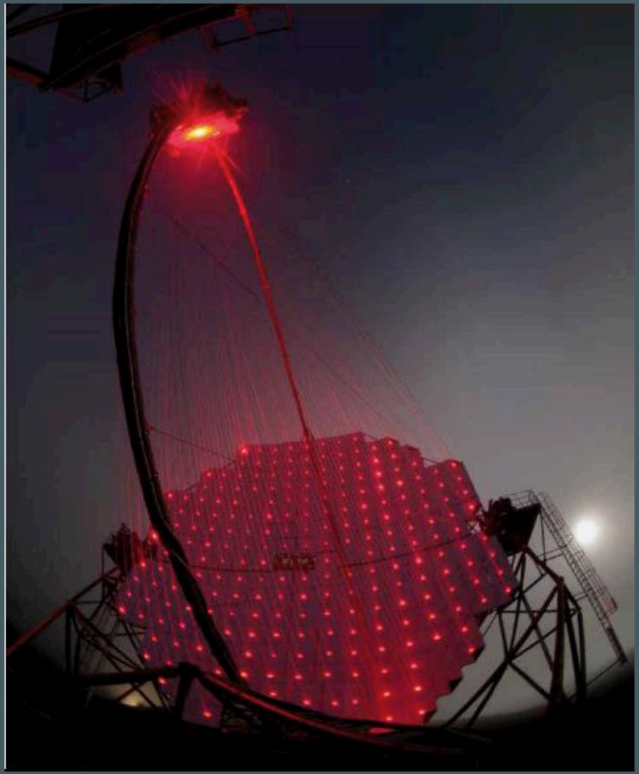


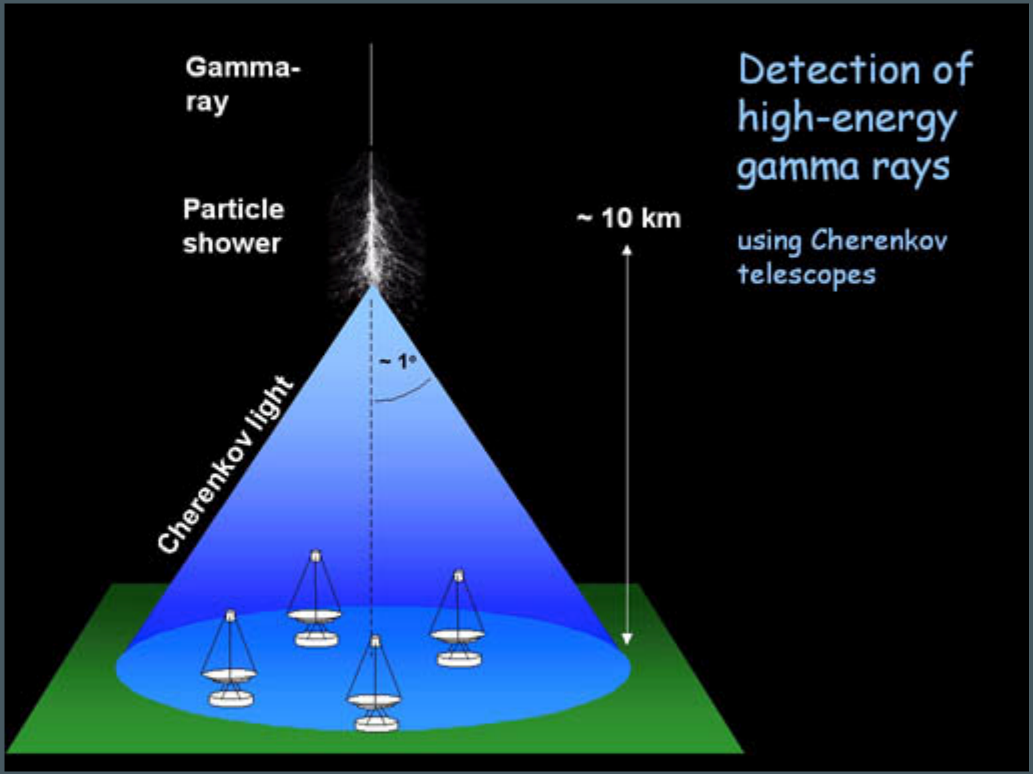


Atmospheric Cherenkov Telescopes

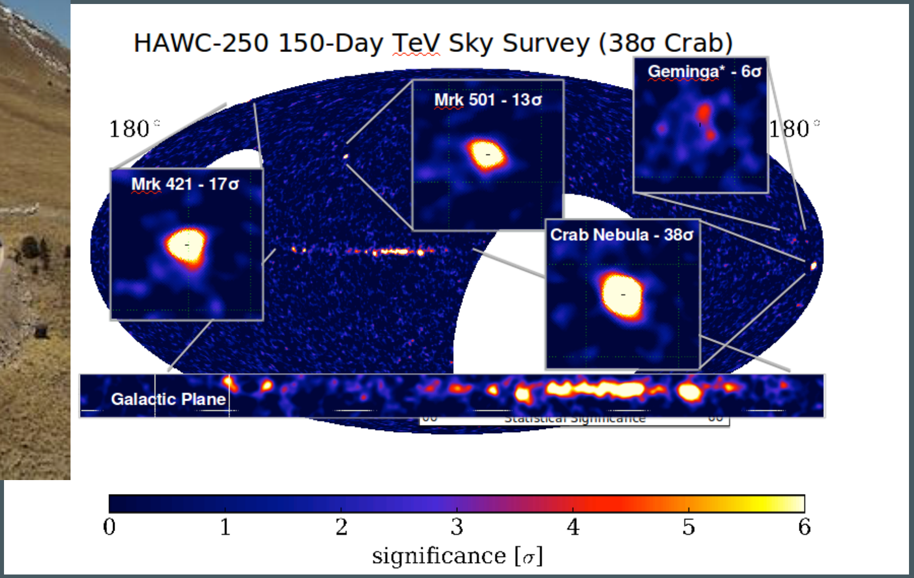


Atmospheric Cherenkov Telescopes



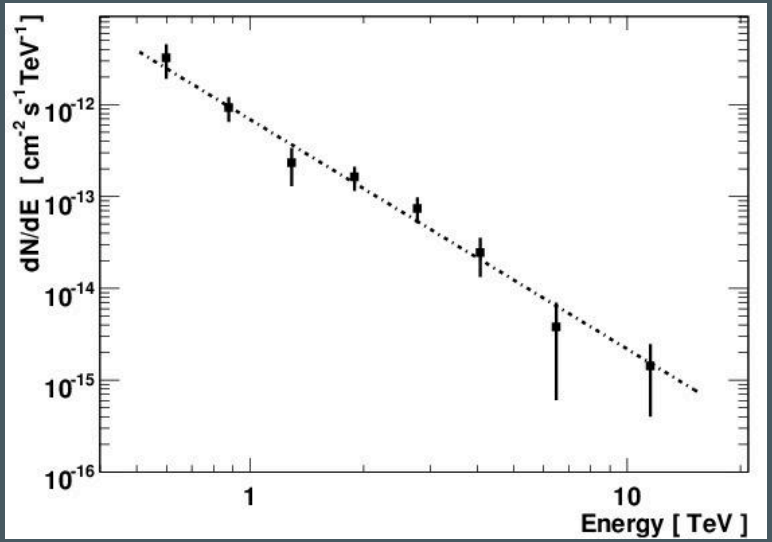


HAWC

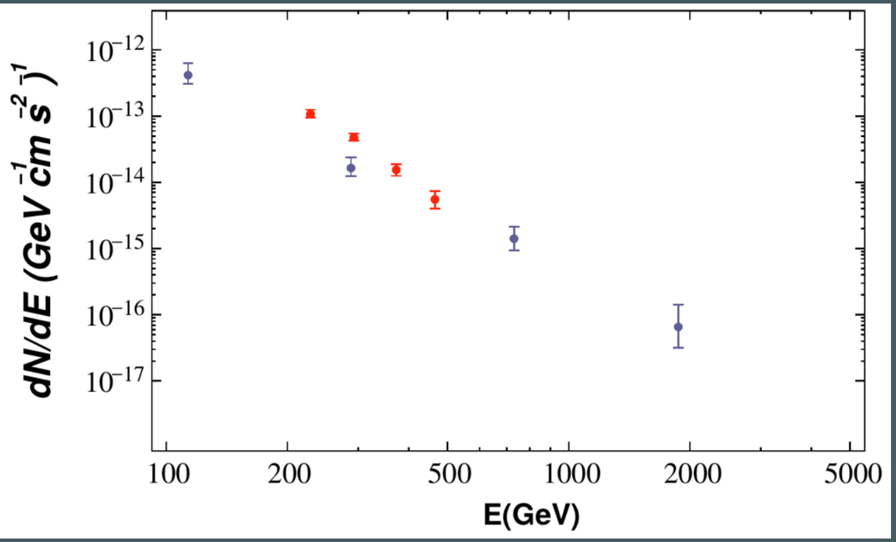

HAWC-250 150-Day TeV Sky Survey (38σ Crab)
Geminga* - 6σ
Mrk 501 - 13σ
180°
180°
Mrk 421 - 17σ
Crab Nebula - 38σ
Galactic Plane
0
1
2
3
4
5
6
significance [σ]

Fermi gamma-ray space telescope

HESS(black), MAGIC (blue), VERITAS (red)

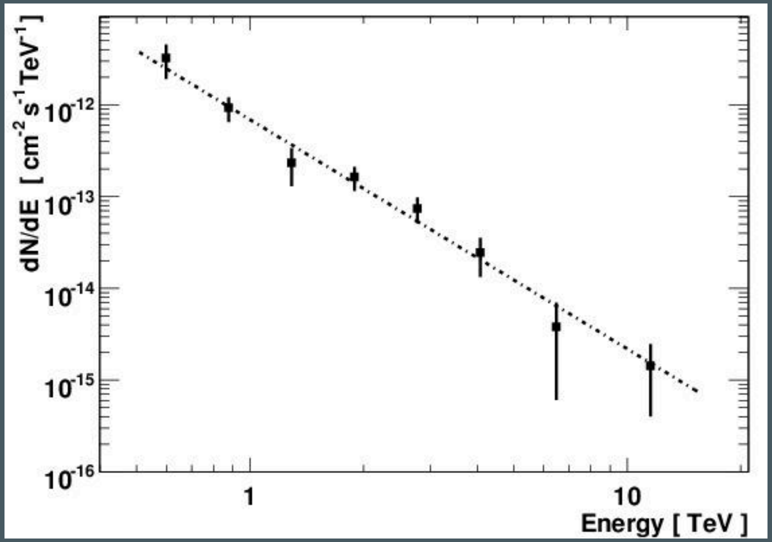


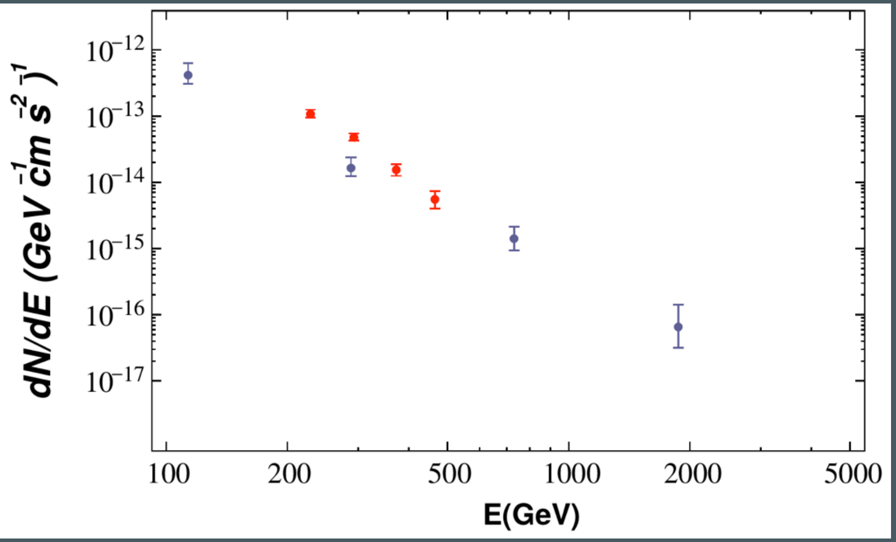
1 ES0229+200 (z=0.14)



3C66A (z=0.44)

HESS(black), MAGIC (blue), VERITAS (red)



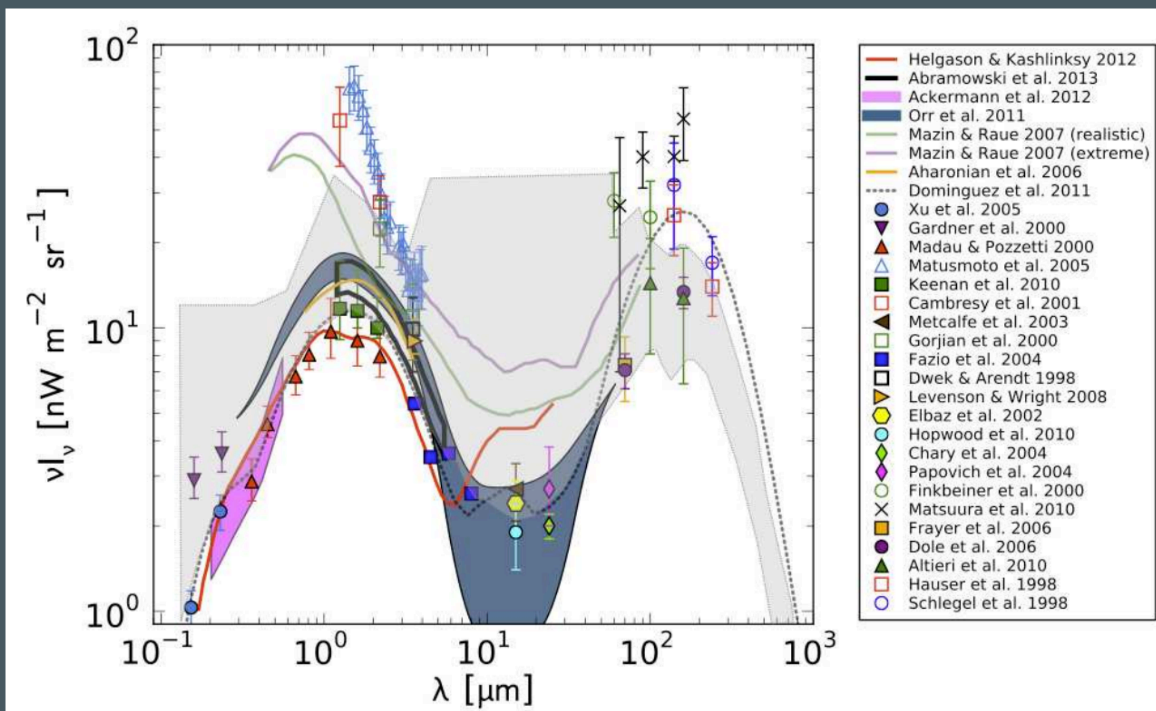


1 ES0229+200 (z=0.14) 3C66A (z=0.44)

Theory: "we predict a sharp cutoff between 0.1 and 1 TeV" Stecker, et al. (1992)

Data: no sign of absorption due to $\gamma\gamma_{EBL} \rightarrow e^+e^-$

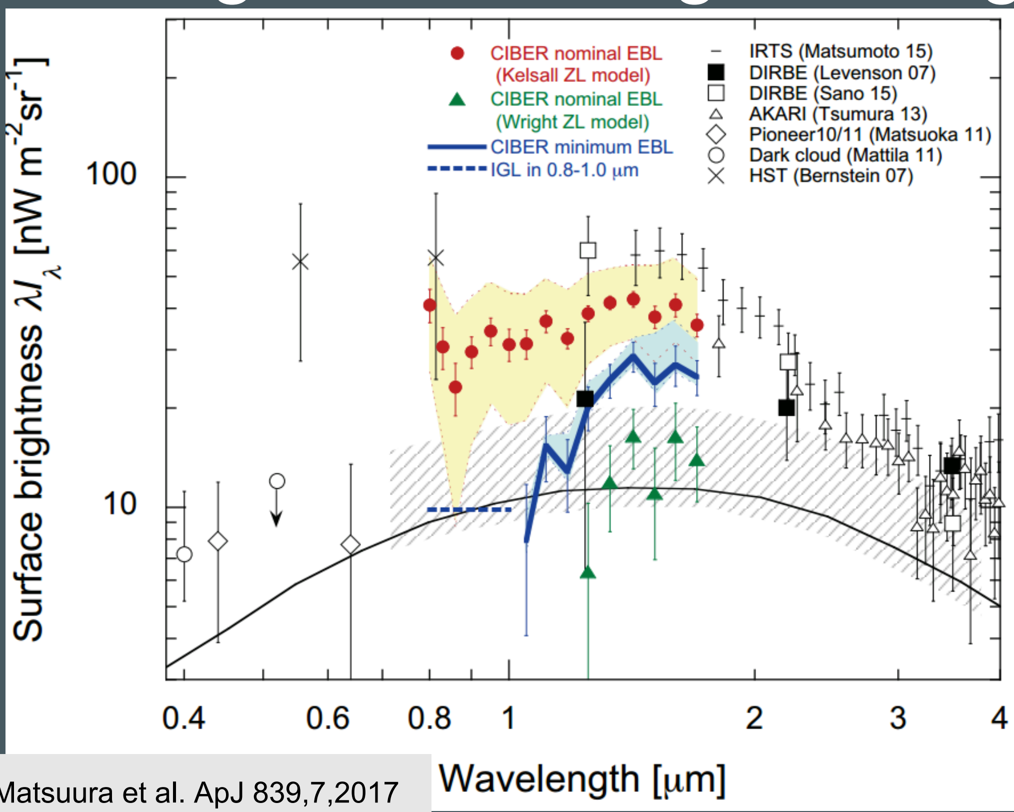
Extragalactic background light



InfraNed:
Ned Wright's image in nfrared light. Credits: NASA/WISE

Interactions with EBL must degrade the energies of TeV photons: $\gamma\gamma_{EBL} \rightarrow e^+e^-$

Extragalactic background light

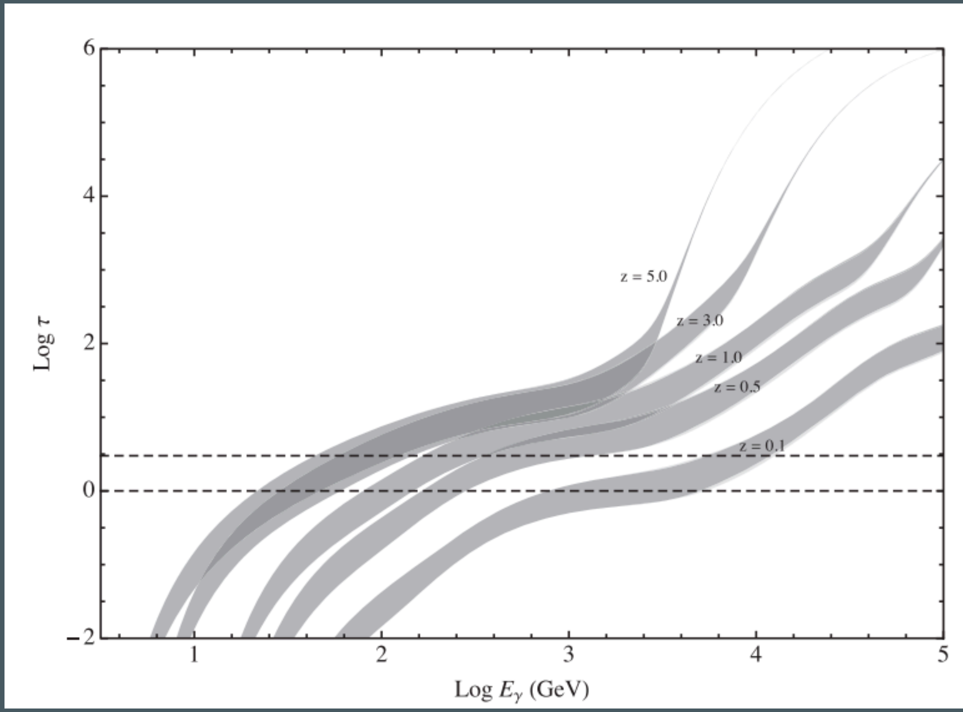


Matsuura et al. ApJ 839,7,2017

InfraNed:
Ned Wright's image in nfrared light. Credits: NASA/WISE

Interactions with EBL must degrade the energies of TeV photons: $\gamma\gamma_{EBL} \rightarrow e^+e^-$

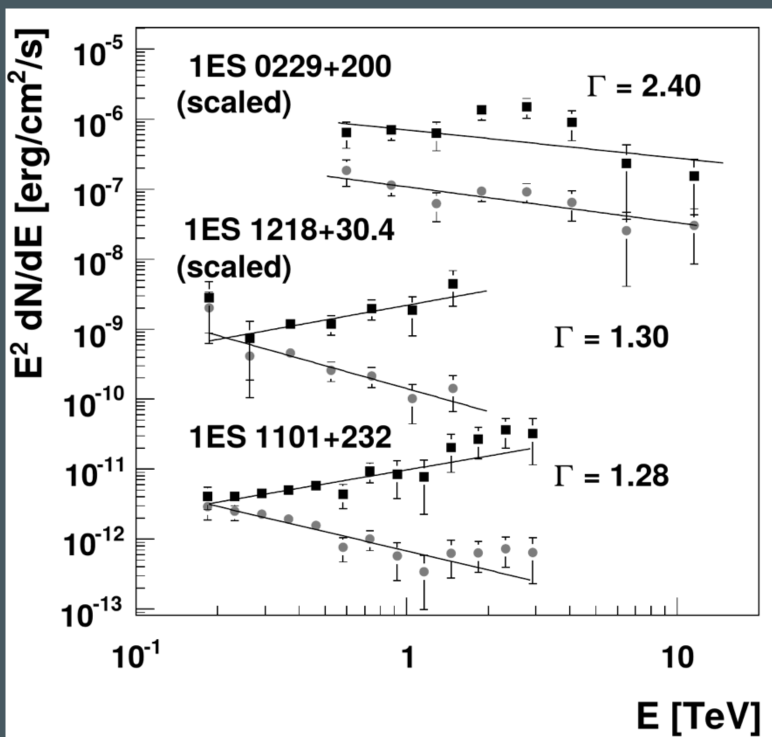
Optical depth to gamma rays



Strong suppression of the gamma ray spectrum expected for E > TeV, z > 0.3

[Stecker, Scully, Malkan, 2016]

Distant blazars: implausibly hard spectra?

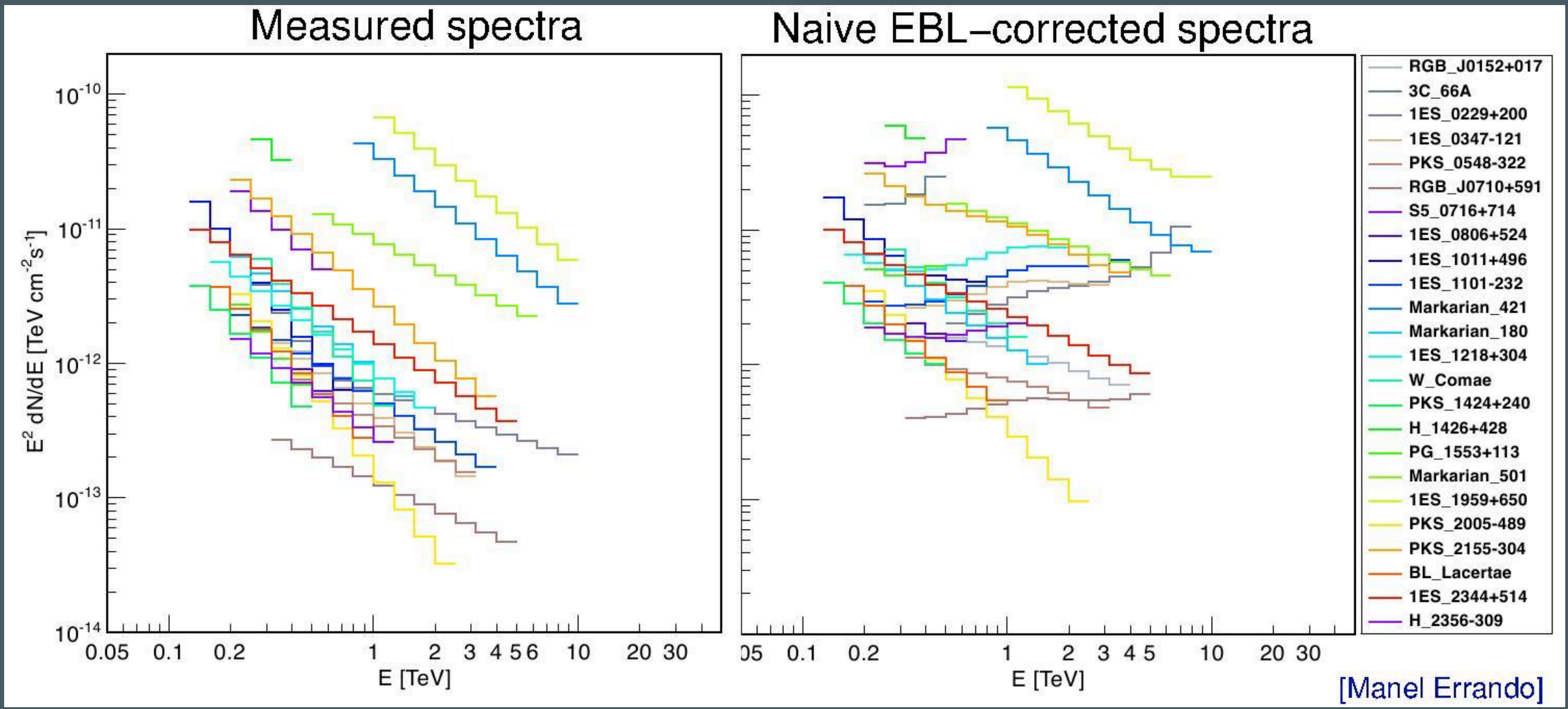


Absorption-corrected spectra would have to be extremely hard for distant blazars:

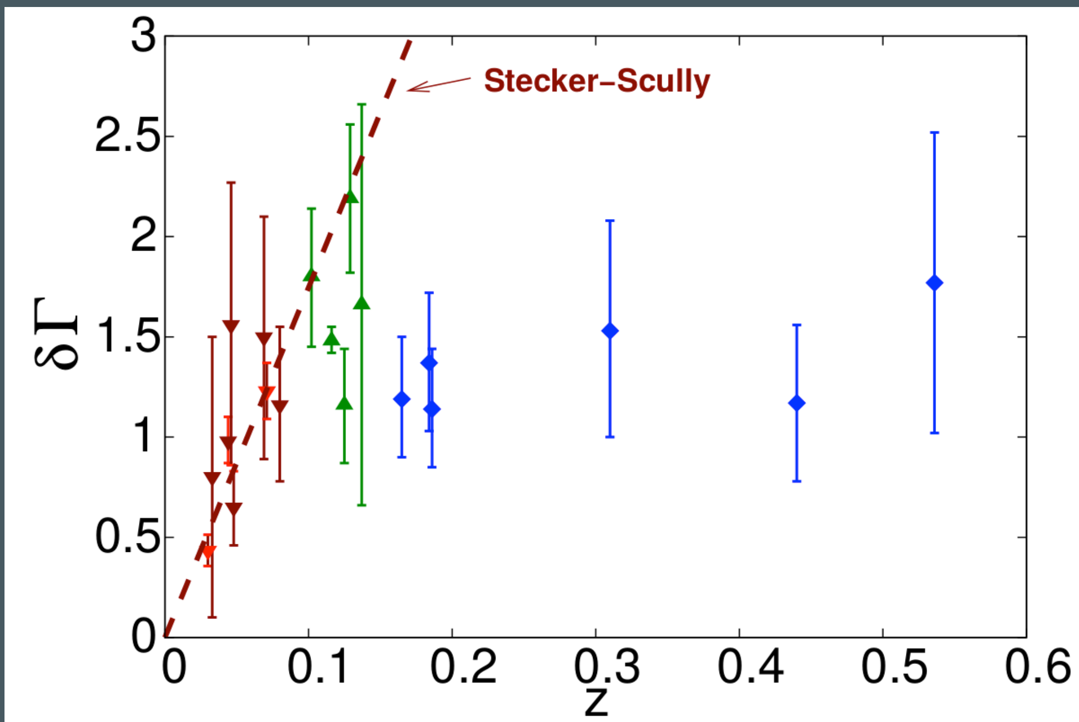
$$\Gamma < 1.5$$

[Aharonian et al.]

Blazar spectra



Spectral softening: problem with distant blazars



$\delta\Gamma = \Gamma_{GeV} - \Gamma_{TeV}$ [Stecker, Scully]

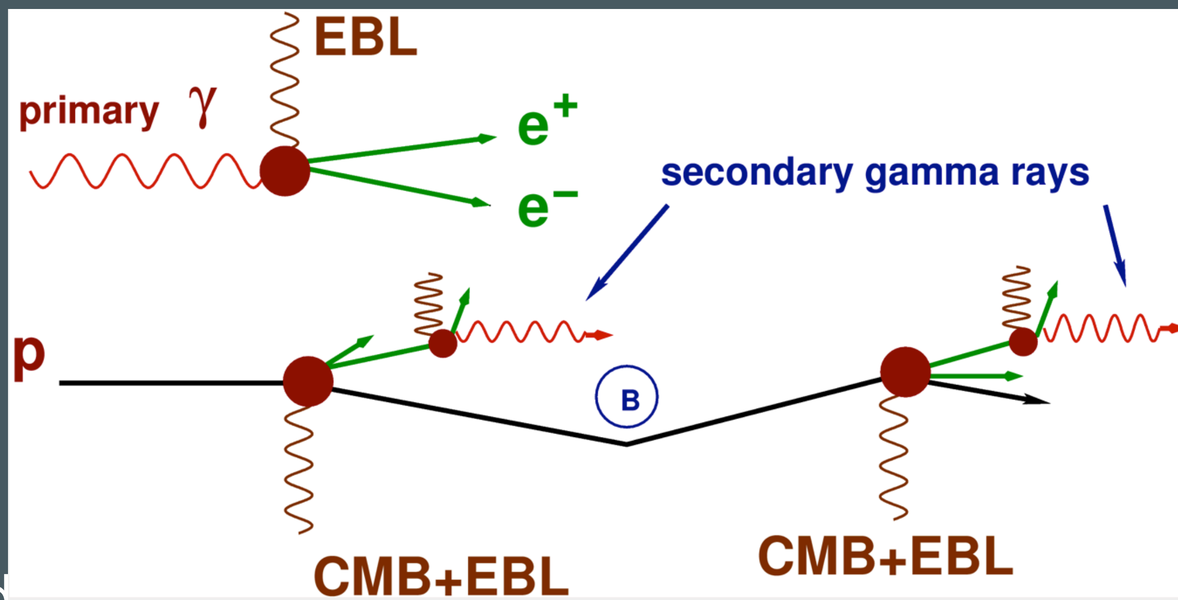
Analytical predictions for the spectral softening work well for the nearby blazars, but not for distant blazars

The mysterious transparency of the Universe...

- Hypothetical axion-like particles: photons convert into them in magnetic fields near the source, and they convert back to gamma rays? [de Angelis et al.]
- Violation of ~~$\gamma\gamma_{EBL} \rightarrow e^+e^-$~~ iance suppresses the pair production? [Stecker, Glashow]

New physics is an exciting possibility,
but can there be a more conventional explanation?

γ rays and cosmic rays



Secondary gamma rays from line-of-sight interactions of CRs

[Essey & AK (2010)]

Different scaling

$$F_{\mathrm{primary},\gamma}(d) \propto \frac{1}{d^2}\exp\{-d/\lambda_\gamma\}$$

$$F_{\mathrm{secondary},\gamma}(d) = \frac{p\lambda_\gamma}{4\pi d^2}\left[1 - e^{-d/\lambda_\gamma}\right] \propto \begin{cases} 1/d, & \text{for } d \ll \lambda_\gamma, \\ 1/d^2, & \text{for } d \gg \lambda_\gamma. \end{cases}$$

$$F_{\mathrm{secondary},\nu}(d) \propto (F_{\mathrm{protons}} \times d) \propto \frac{1}{d}.$$

For distant sources, the secondary signal wins!

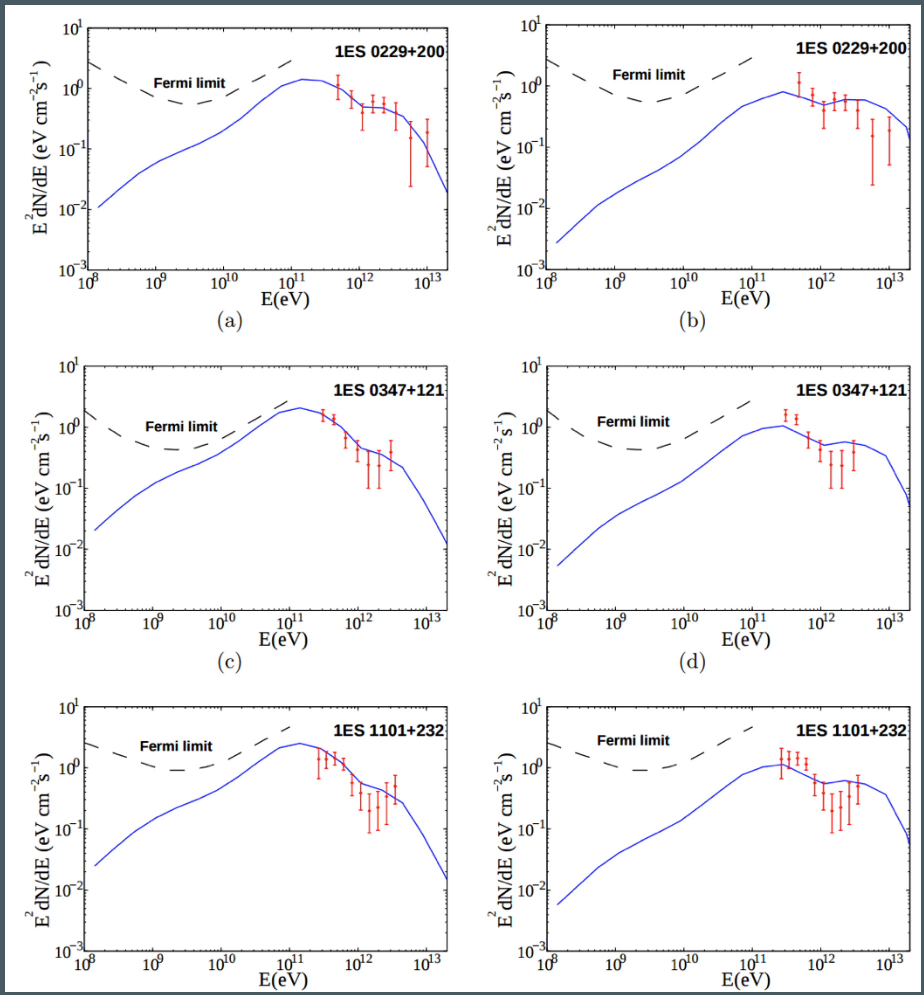
One-parameter fit (power in CR) for each source
[Essey & AK (2010); Essey, Kalashev, AK, Beacom (2011)]

Good agreement with data for high-redshift blazars
(both "high" and "low" EBL models).

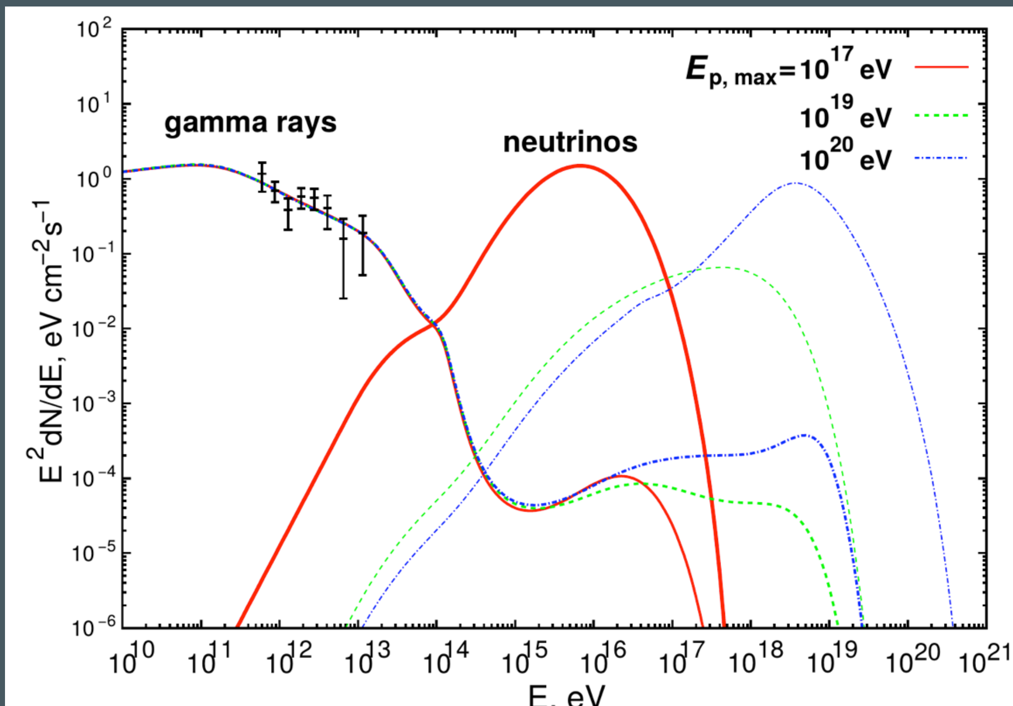
Reasonable CR power for a source up to z~1
[Aharonian, Essey, AK, Prosekin (2013);
Razzaque, Dermer, Finke (2012);
Murase, Dermer, Takami, Migliore (2012)]

Consistent with data on time variability
[Prosekin, Essey, AK, Aharonian (2012)]

Essey, Kalashev, AK, Beacom, ApJ (2011)



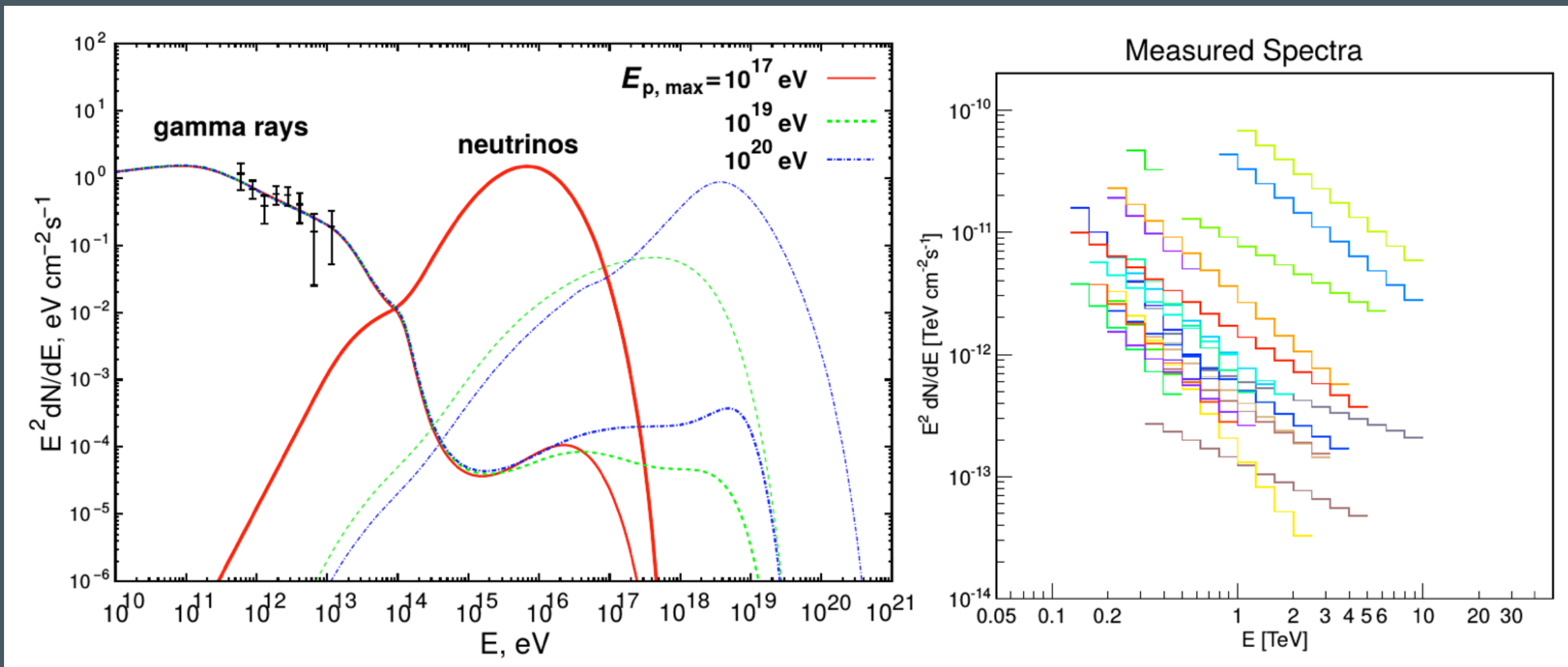
Secondary γ,ν from 1ES0229+200 (z=0.14)



- Gamma-ray spectra **robust**
- Neutrino spectra **peaked**

[Essey, Kalshev, AK, Beacom, PRL (2010)]

Robust shapes explain observed universality



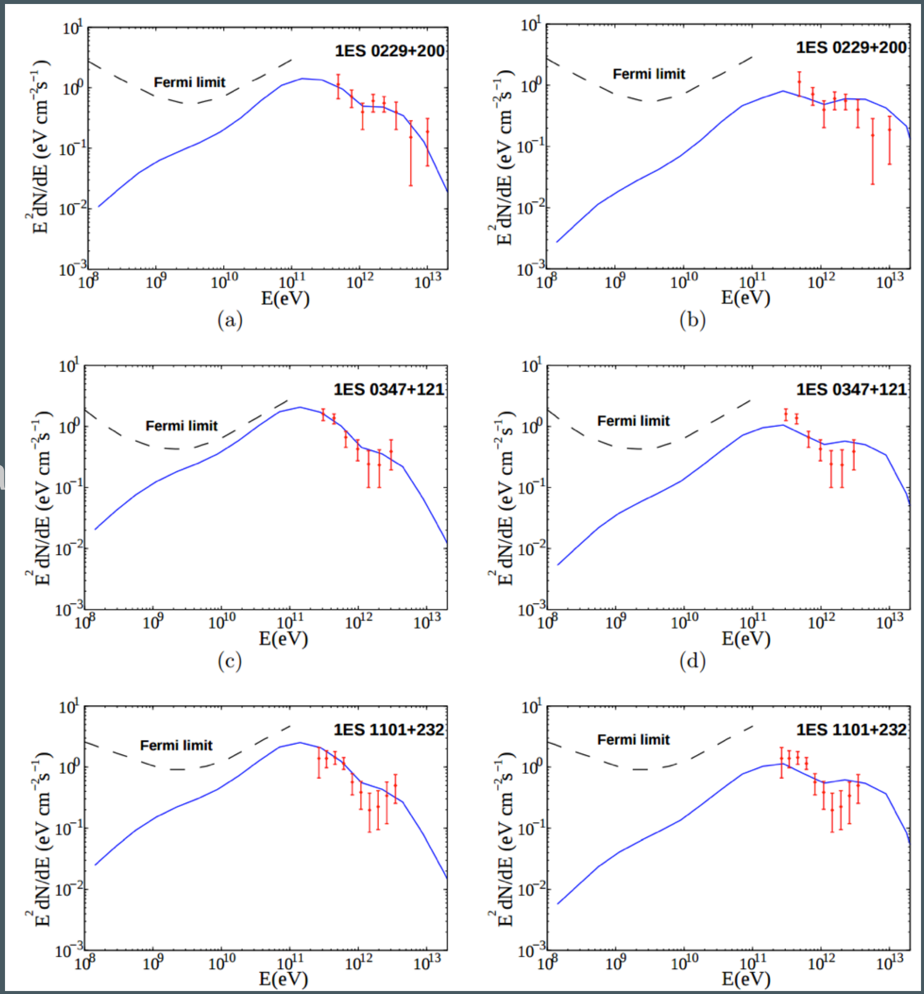
EBL models

“Low EBL” on the left,

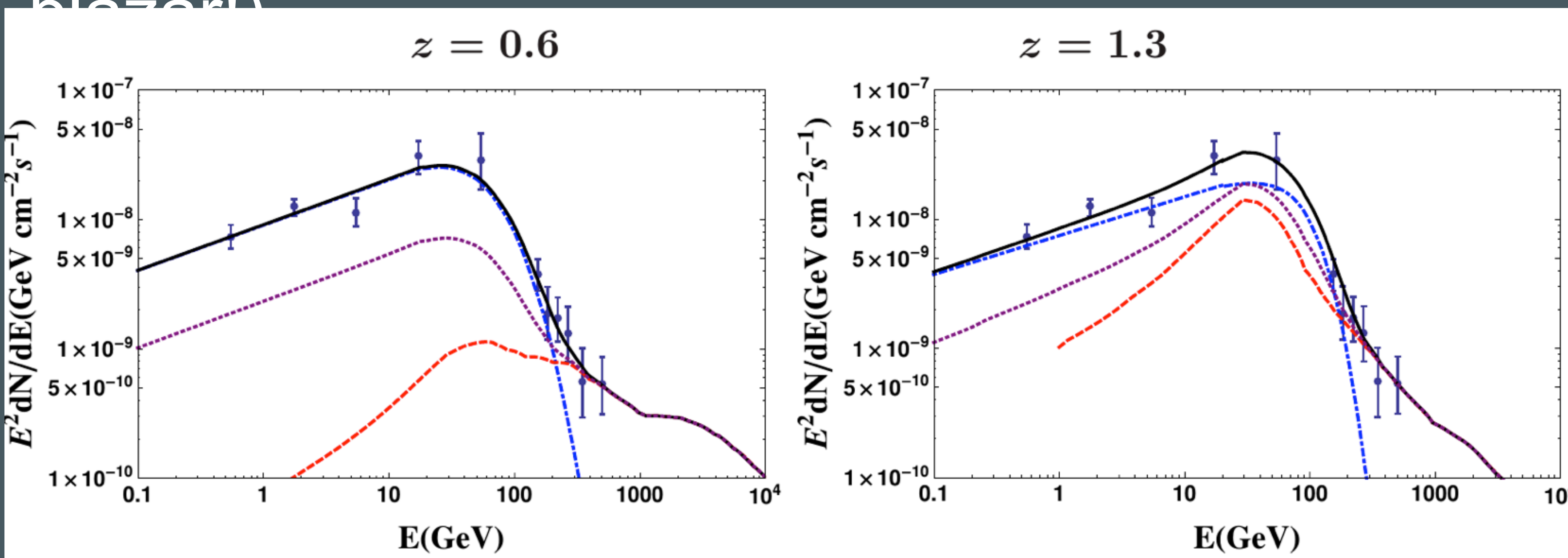
“High EBL” on the right,

Both appear to be consistent. More data needed to distinguish..

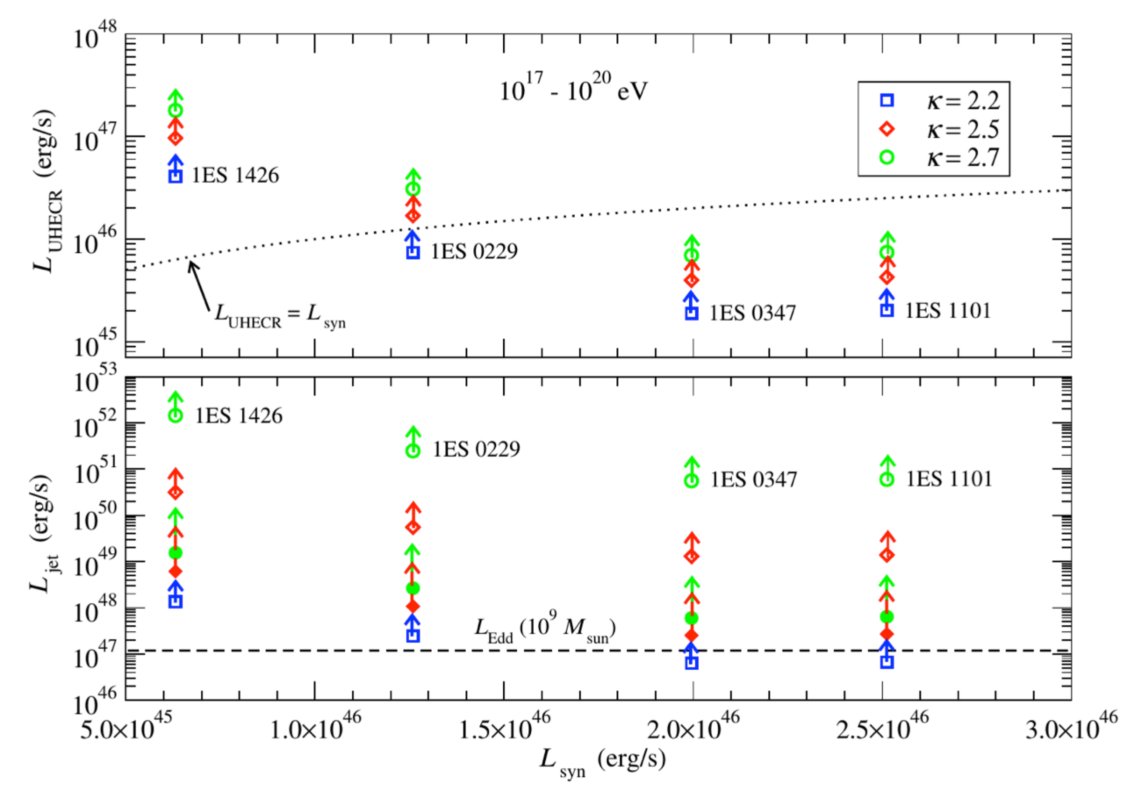
Source	Redshift	EBL Model	L_p , erg/s	$L_{p,\rm iso}$, erg/s	χ^2	DOF
1ES0229+200	0.14	Low	1.3×10^{43}	4.9×10^{45}	6.4	7
1ES0229+200	0.14	High	3.1×10^{43}	1.1×10^{46}	1.8	7
1ES0347-121	0.188	Low	2.7×10^{43}	1.0×10^{46}	16.1	6
1ES0347-121	0.188	High	5.2×10^{43}	1.9×10^{46}	3.4	6
1ES1101-232	0.186	Low	3.0×10^{43}	1.1×10^{46}	16.1	9
1ES1101-232	0.186	High	6.3×10^{43}	2.3×10^{46}	4.9	9



PKS 1424+240 at z>0.6 (the most extreme TeV blazar!)



Required power in cosmic rays

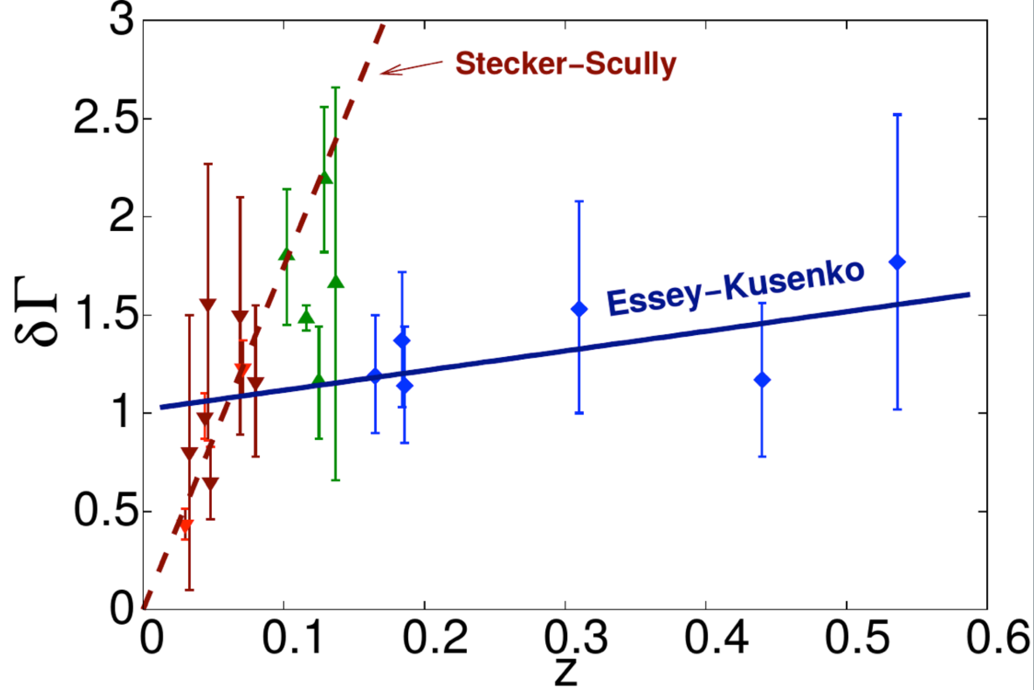


High, but not unreasonable

Consistent with models

[Razzaque et al. (2012)]

Spectral softening



$$\delta\Gamma = \Gamma_{\mathrm{GeV}} - \Gamma_{\mathrm{TeV}}$$

Three populations in red, blue and green are seen in primary, secondary, or mixed components, respectively.

Predictions: no variability for TeV blazars at z>0.15. In good agreement with data. [Prosekin, Essey, AK, Aharonian]

CTA extragalactic survey discovery potential

Cherenkov Telescope Array (CTA)
extragalactic survey will see an enhancement in the number of
distant TeV sources, thanks to secondary gamma rays.

[De Franco, Inoue, Sanchez-Conde, Cotter (2017)]

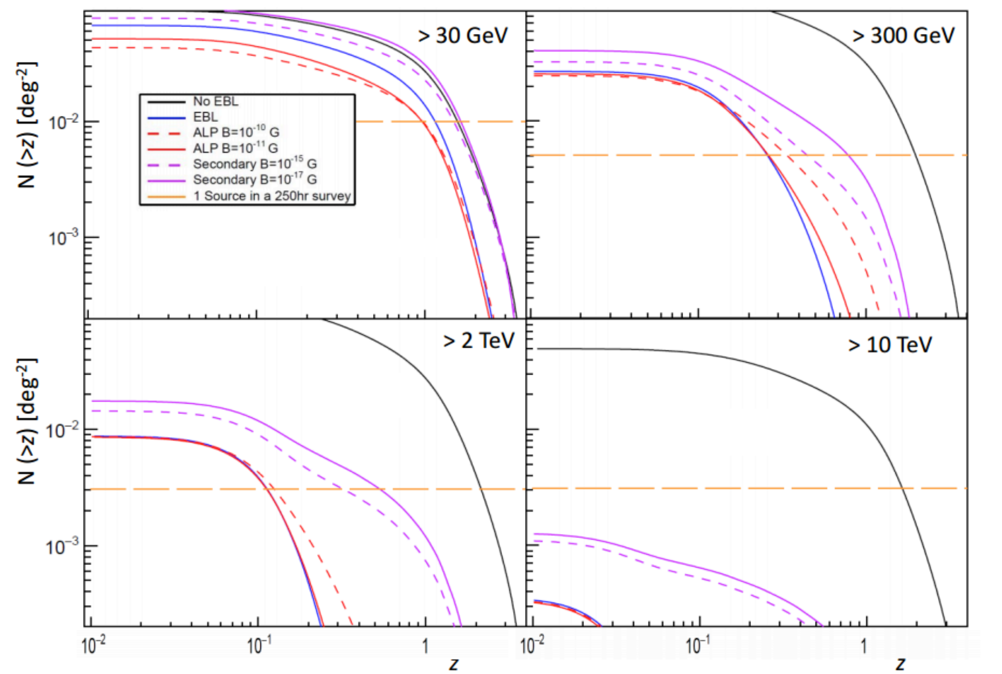


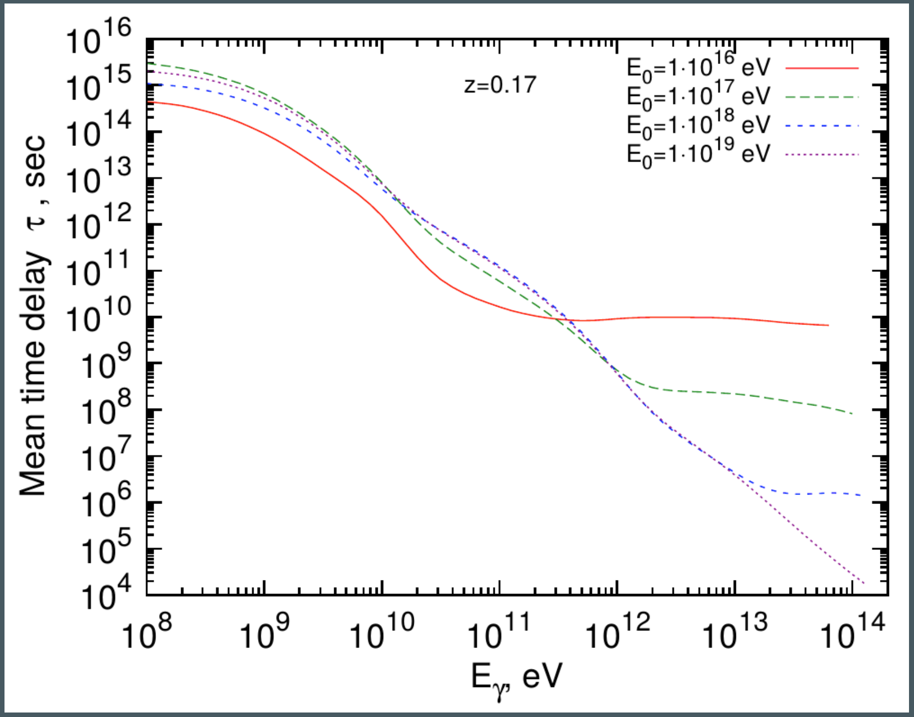
Figure 3: Cumulative source count as a function of redshift for CTA-South, assuming a 5σ integral flux sensitivity and 50 hr exposure observation. Dashed horizontal line represents 1 source detected in a 250 hr survey.

Erosion of time variability for *E>1 TeV, z>0.15*

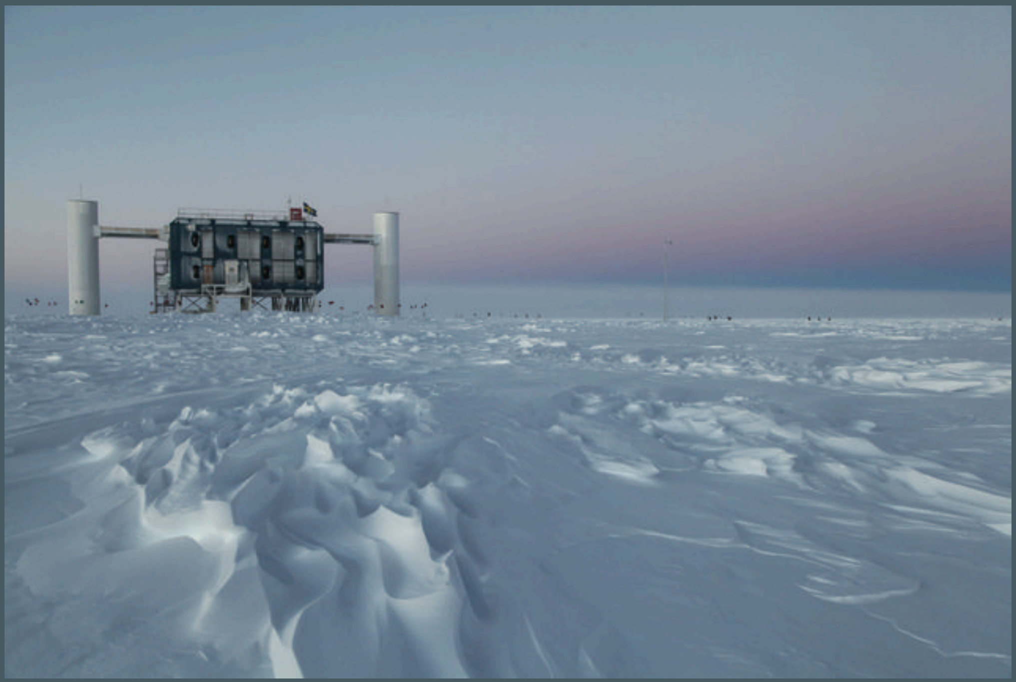
Nearby blazars are variable at all energies. Distant blazars are variable at lower energies, but there is no evidence of variability for, e.g., E>1 TeV, z > 0.15

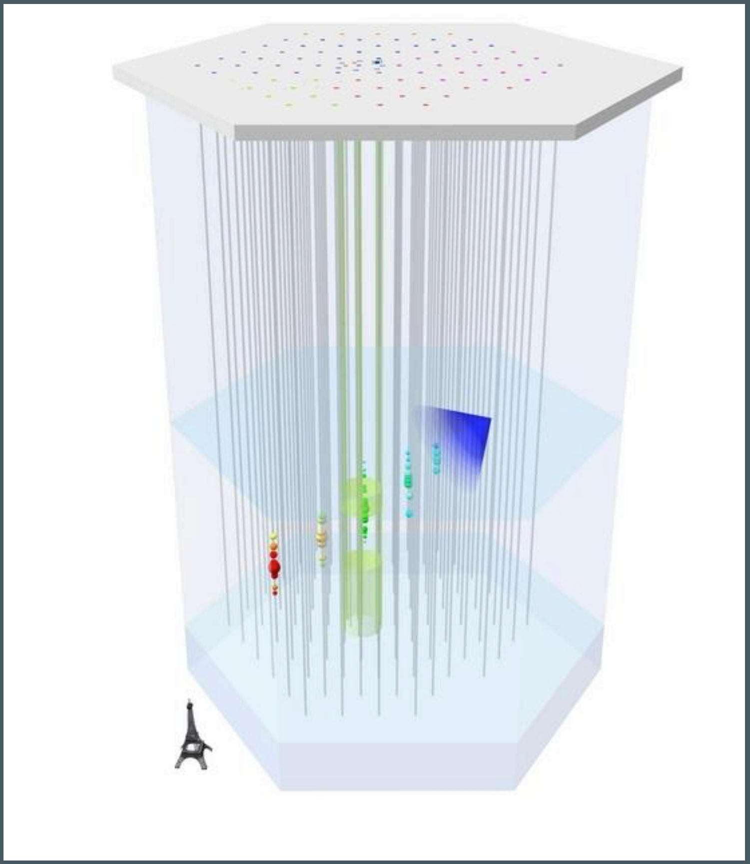
Prediction: stochastic *pedestal* emerges at high energy, high redshifts, for distant blazars above which some flares may rise in a stochastic fashion.

[Prosekin, Essey, AK, Aharonian, ApJ 757 (2012) 183]



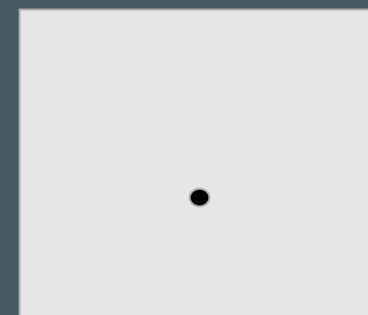
IceCube detector





IceCube detector

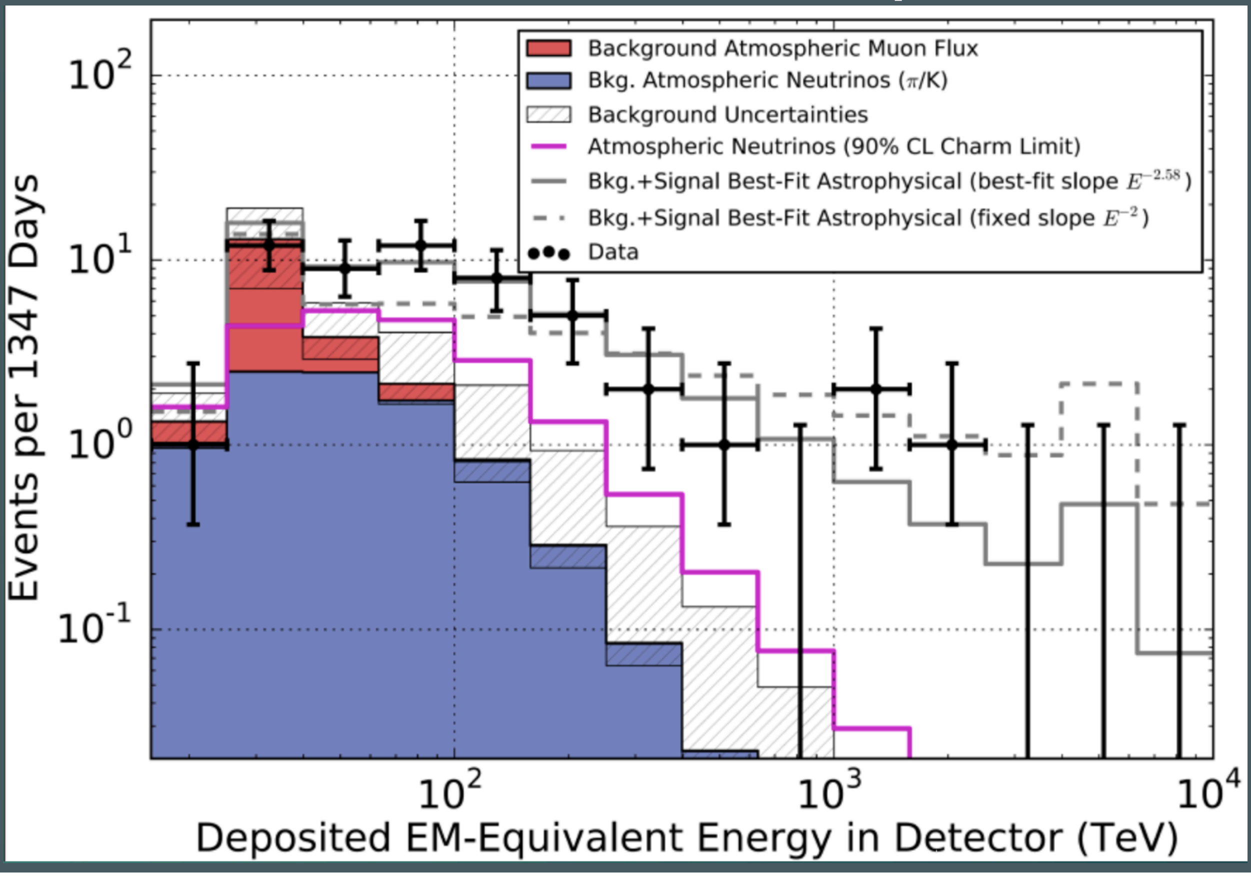
NeutrinoNed: image of Ned in νs



Emits 400 million neutrinos per day, mainly from about 20-25 mg of ^{40}K

Absorbed only 1 neutrino in his his first 70 years (out of 10^{23} solar neutrinos that passed through.)

IceCube neutrinos: the spectrum

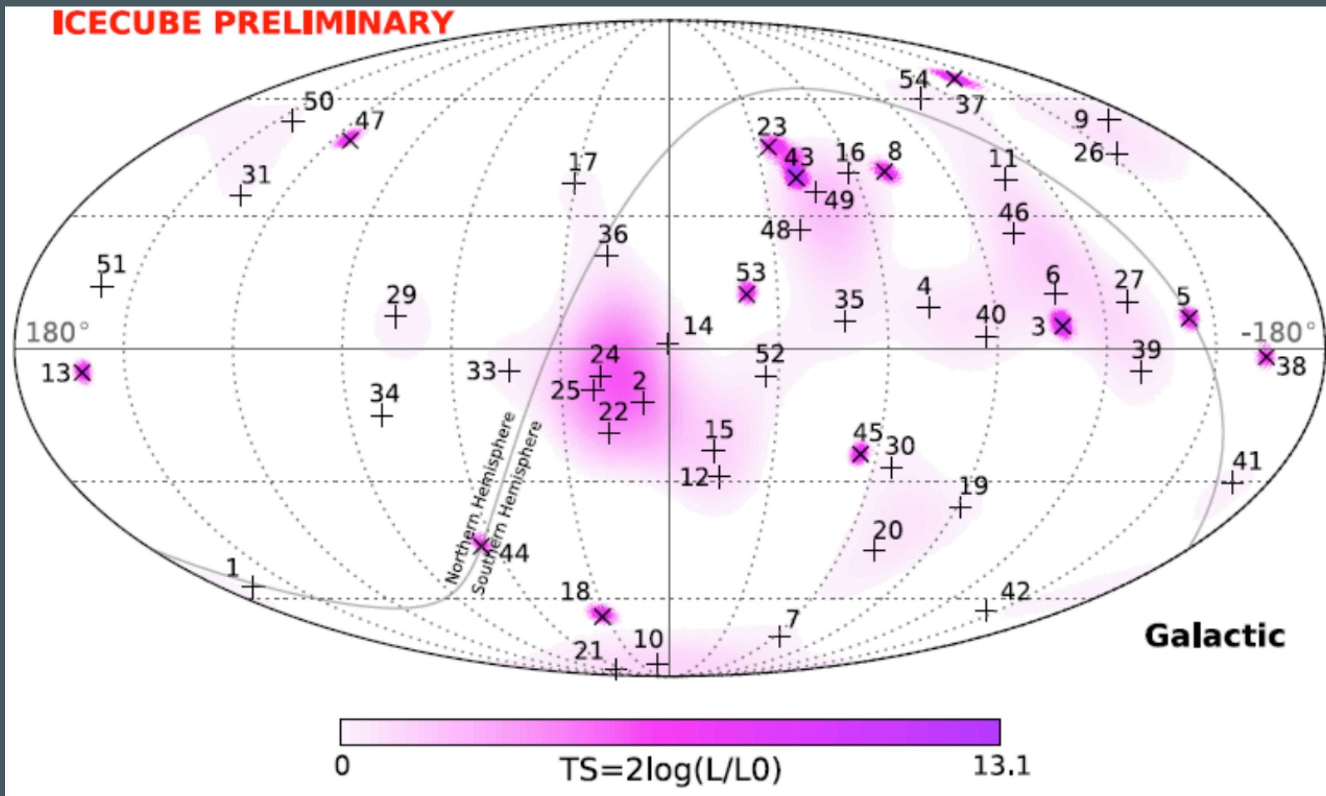


Power law with a cutoff?

Two components?

A peak at 1 PeV?

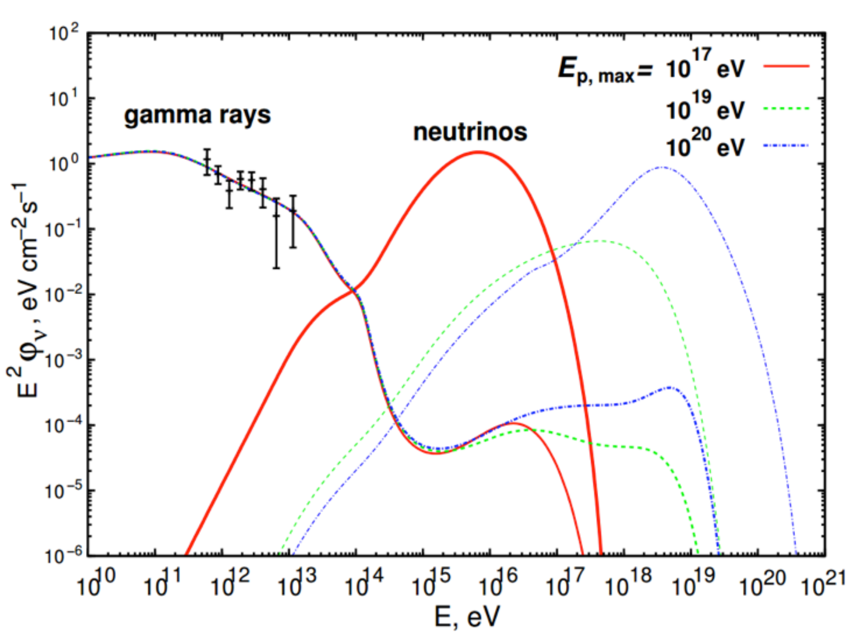
IceCube neutrinos: the arrival directions

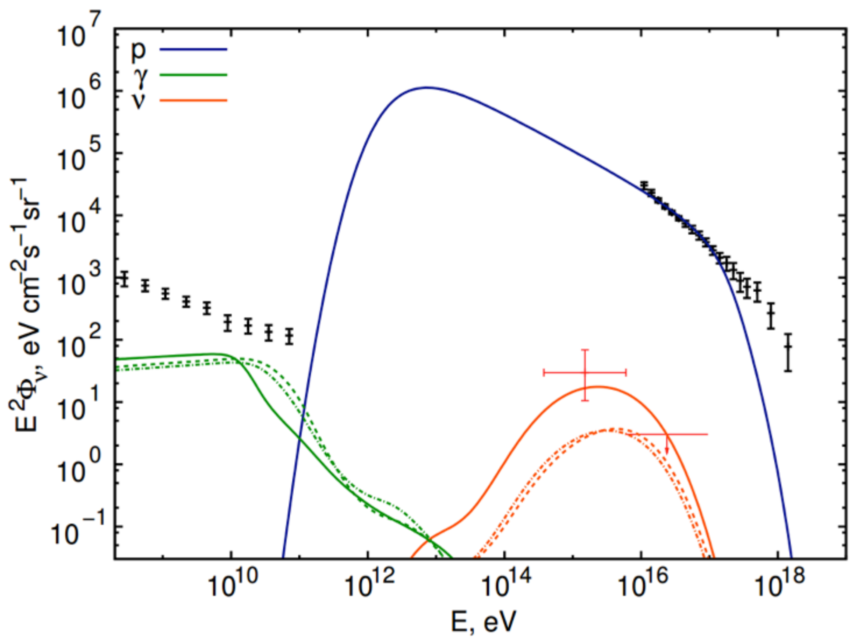


Arrival directions do not appear to trace nearby matter distribution ⇒

Consistent with production on intervening backgrounds, not in sources.

Line-of-sight interactions of CRs from blazars





Essey et al. Phys.Rev.Lett. 104 (2010) 141102; Kalashev et al., Phys.Rev.Lett. 111 (2013) 4, 041103

A peaked spectrum at 1 PeV can result from cosmic rays accelerated in AGN and interacting with photon backgrounds, assuming that secondary photons explain the observations of TeV blazars.

prediction: PRL 104, 141102 (2010)
consistency with IceCube: PRL 111, 041103 (2013)

PRL **104**, 141102 (2010) PHYSICAL REVIEW LETTERS week ending 9 APRIL 2010

Secondary Photons and Neutrinos from Cosmic Rays Produced by Distant Blazars

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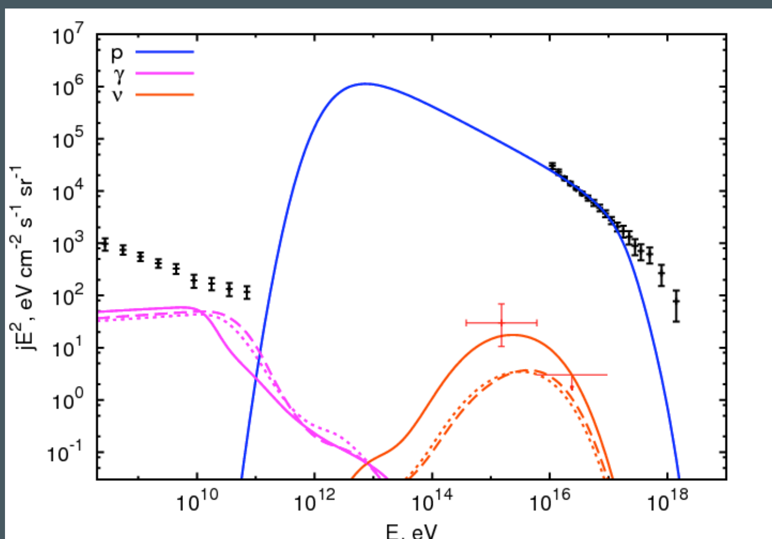
(Received 27 December 2009; revised manuscript received 22 February 2010; published 8 April 2010)

Secondary photons and neutrinos produced in the interactions of cosmic ray protons emitted by distant active galactic nuclei (AGN) with the photon background along the line of sight can reveal a wealth of new information about the intergalactic magnetic fields, extragalactic background light, and the acceleration mechanisms of cosmic rays. The secondary photons may have already been observed by gamma-ray telescopes. We show that the secondary neutrinos improve the prospects of discovering distant blazars by IceCube, and we discuss the ramifications for the cosmic backgrounds, magnetic fields, and AGN models.

DOI: 10.1103/PhysRevLett.104.141102 PACS numbers: 95.85.Pw, 98.54.Cm, 98.70.Sa, 95.85.Ry



PRL **111**, 041103 (2013) PHYSICAL REVIEW LETTERS week ending 26 JULY 2013

PeV Neutrinos from Intergalactic Interactions of Cosmic Rays Emitted by Active Galactic Nuclei

Oleg E. Kalashev,[1] Alexander Kusenko,[2,3] and Warren Essey[2]
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(Received 28 February 2013; revised manuscript received 14 June 2013; published 24 July 2013)

The observed very high energy spectra of *distant* blazars are well described by secondary gamma rays produced in line-of-sight interactions of cosmic rays with background photons. In the absence of the cosmic-ray contribution, one would not expect to observe very hard spectra from distant sources, but the cosmic ray interactions generate very high energy gamma rays relatively close to the observer, and they are not attenuated significantly. The same interactions of cosmic rays are expected to produce a flux of neutrinos with energies peaked around 1 PeV. We show that the diffuse isotropic neutrino background from many distant sources can be consistent with the neutrino events recently detected by the IceCube experiment. We also find that the flux from any individual nearby source is insufficient to account for these events. The narrow spectrum around 1 PeV implies that some active galactic nuclei can accelerate protons to EeV energies.

DOI: 10.1103/PhysRevLett.111.041103 PACS numbers: 95.85.Ry, 98.54.Cm, 98.70.Sa

Implications for intergalactic magnetic fields

Magnetic fields along the line of sight:

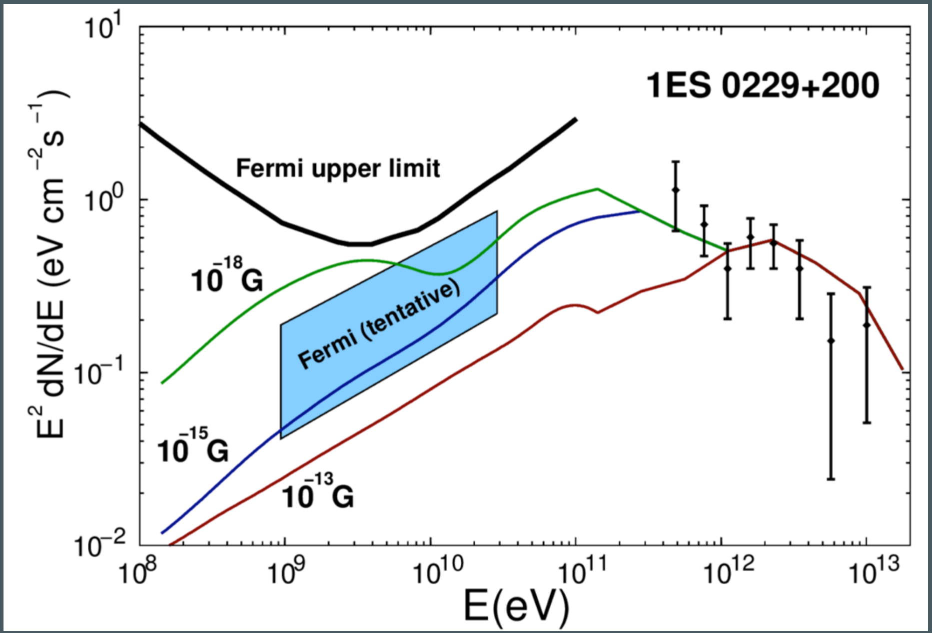
$$1 \times 10^{-17}\ \mathrm{G} < B < 3 \times 10^{-14}\ \mathrm{G}$$

Essey, Ando, AK (2011)

Lower limits: see also Finke et al. (2015)

If an intervening filament deflects protons, then no secondary component is expected.

However, even a source at z~1 has an order-one probability to be unobscured by magnetic fields, and can be seen in secondary gamma rays
[Aharonian, Essey, AK, Prosekin, arXiv: 1206.6715]



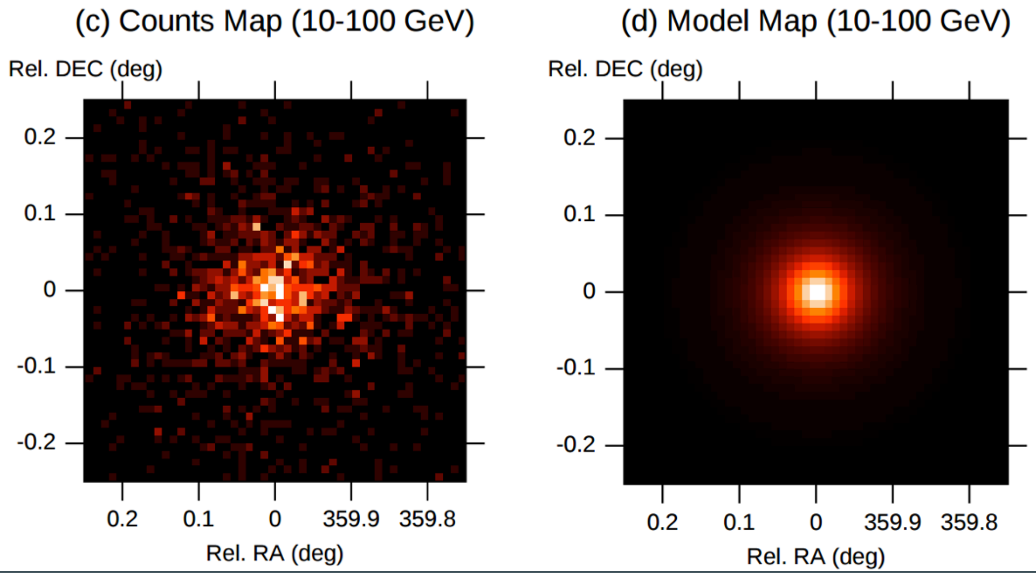
Essey, Ando, AK (2011)

Blazar halos: an independent measurement of IGMFs

Halos around stacked images of blazars implying

B~10^{-15} G

were reported (3.5σ) in 1st year Fermi data [Ando & AK, ApJL 722 (2010) L39].



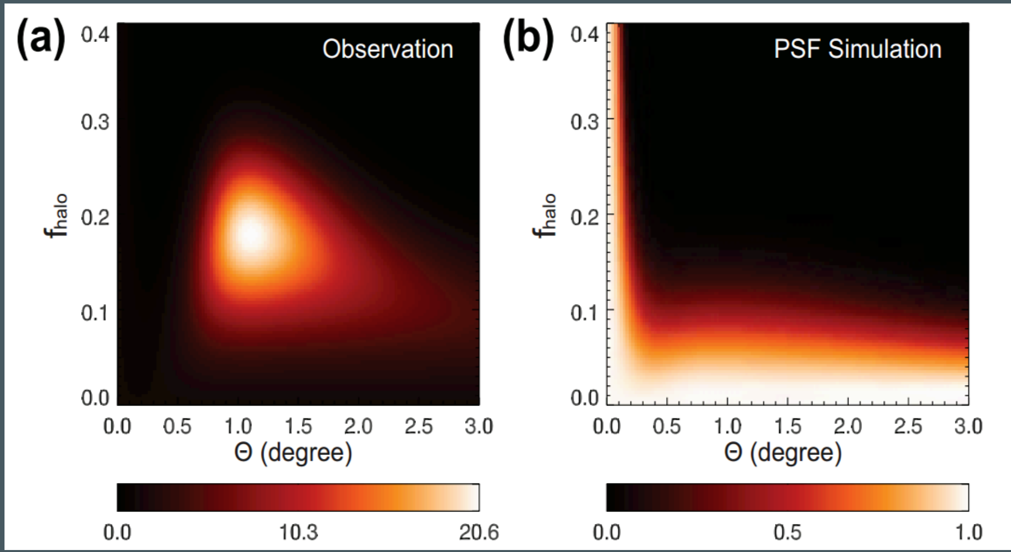
Ando & AK, ApJL 722 (2010)

Blazar halos: an independent measurement of IGMFs

Halos around stacked images of blazars implying **B~10^{-15} G** were reported (3.5σ) in 1st year Fermi data [Ando & AK, ApJL 722 (2010) L39].

Now the same technique was applied to the much larger Fermi data set, detecting lower energy halos of z< 0.5 blazars. The results, **B~10^{-17} -- 10^{-15} G** [Chen, et al. (2015)], confirm earlier results of Ando & AK, arXiv: 1005.1924.

Consistent with independent measurement based on the gamma-ray spectra of blazars [Essey, Ando, AK, arXiv:1012.5313]



Chen, Buckley, Ferrer, Phys. Rev. Lett. (2015) confirm halos, IGMFs in the **B~10^{-17} -- 10^{-15} G** range

Extragalactic magnetic fields: a new window on the early

Conclusion

- We have learned a lot from treating gamma rays and cosmic rays consistently
- Excellent agreement of gamma-ray spectra with observations of distant blazars (and very little model dependence)
- Neutrinos are an interesting probe (but predictions are model-dependent)
- Now as we understand the "beam", we can use it to test the cosmic photon backgrounds (EBL) and magnetic fields
- The first measurements of the magnetic fields are exciting:
  possibly, a new window on the universe

Happy birthday, Ned!